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	ID	σ8	Statistic	Field	$m_{ m lim}$	CosVar	E/B	Z,	Г
	[18] Maoli et al. 01	$1.03\pm0.05$	$\langle \gamma^2 \rangle$	VLT+CTIO+ WHT+CFHT	82	no	no	2	0.21
	[27] LVW et al. 01	$0.88\pm0.11$	$\begin{array}{c} \langle \gamma^2 \rangle,  \xi(r) \\ \langle M^2_{\mathrm{ap}} \rangle \end{array}$	CFHT 8 sq.deg.	I=24	no	no (yes)	1.1	0.21
	[22] Rhodes et al. 01	$0.91\substack{+0.25 \\ -0.29}$	$\xi(r)$	HST 0.05 sq.deg.	I=26	yes	no	0.9-1.1	0.25
<u>.</u> კ	[8] Hoekstra et al. 01	$0.81\pm0.08$	$\langle \gamma^2 \rangle$	CFHT+CTIO 24 sq.deg.	R=24	yes	no	0.55	0.21
ں ا	[2] Bacon et al. 02	$0.97\pm0.13$	$\xi(r)$	Keck+WHT 1.6 sq.deg.	R=25	yes	no	0.7-0.9	0.21
G,	[21] Refregier et al. 02	$0.94\pm0.17$	$\langle \gamma^2 \rangle$	HST 0.36 sq.deg.	I=23.5	yes	no	0.8-1.0	0.21
o	[28] LVW et al. 02	$0.94\pm0.12$	$\langle M^2_{\rm ap} \rangle$	CFHT 12 sq.deg.	I=24	yes	yes	0.78-1.08	0.1-0.4
LL_	[9] Hoekstra et al. 02	$0.91\substack{+0.05 \\ -0.12}$	$\begin{array}{c} \langle \gamma^2 \rangle,  \xi(r) \\ \langle M^2_{\rm ap} \rangle \end{array}$	CFHT+CTIO 53 sq.deg.	R=24	yes	yes	0.54-0.66	0.05-0.5
	[4] Brown et al. 02	$0.74\pm0.09$	$\langle \gamma^2  angle,  \xi(r)$	ESO 1.25 sq.deg.	R=25.5	yes	no (yes)	0.8-0.9	-
	[6] Hamana et al. 02	$(2\sigma)0.69^{+0.35}_{-0.25}$	$\langle M^2_{\rm ap} \rangle,  \xi(r)$	Subaru 2.1 sq.deg.	R=26	yes	yes	0.8-1.4	0.1-0.4
	[12] Jarvis et al. 02	$(2\sigma)0.71^{+0.12}_{-0.16}$	$\langle \gamma^2  angle,  \xi(r) \ \langle M^2_{ m ap}  angle$	CTIO 75 sq.deg.	R=23	yes	yes	0.66	0.15-0.5













The	future
Several large cosmic she	ar programs are planned
□ Deep Lens Survey □ CFHT Legacy Survey □ RCS2	28 square degrees (ongoing) 140 square degrees (started) 1000 square degrees (started)
□ LSST □ Pan-STARRS □ SNAP (space)	<ul> <li>&gt; 10<sup>4</sup> square degrees (&gt;2008)</li> <li>&gt; 10<sup>4</sup> square degrees (&gt;2012)</li> <li>few 100 square degrees (&gt;2011)</li> </ul>









Conclusions					
We (might) need to improve our kn	owledge of:				
Source redshift distribution					
photometric redshift from the su	rvey data				
□ deep (photometric) redshift surv	eys				
Non-linear power spectrum					
□ large numerical simulations					
	atric redebifts				
Observational systematic effects					
	5				
□ Detailed simulations					